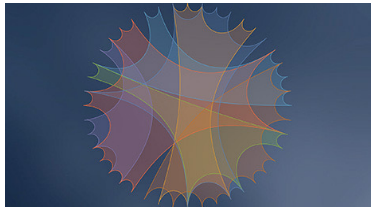
A Universe in Heidelberg

Based on:

[Nature 611, 260-264]

[PRA 106, 033313]

[PRD 105, 105020]



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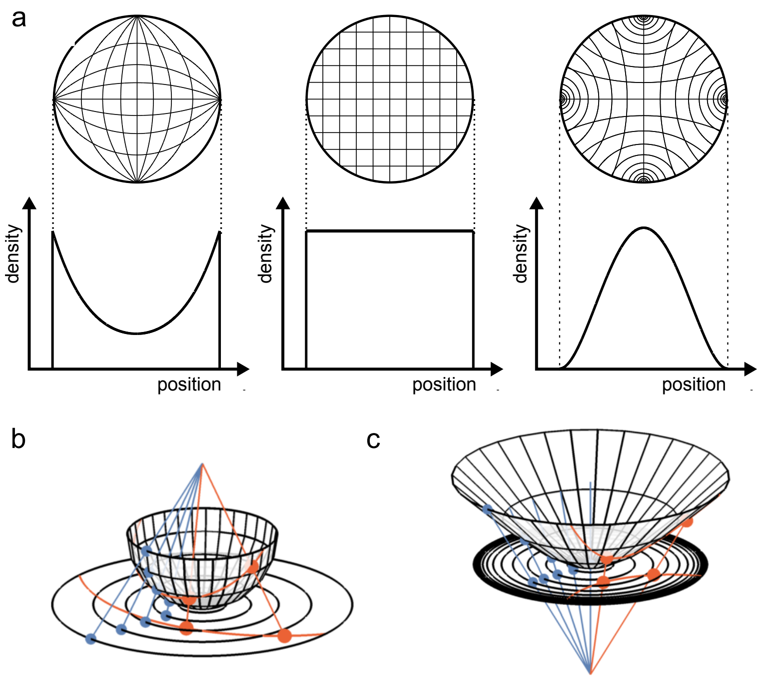

The Team

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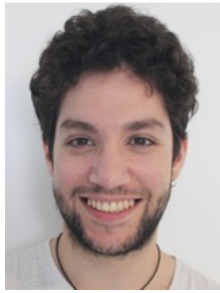
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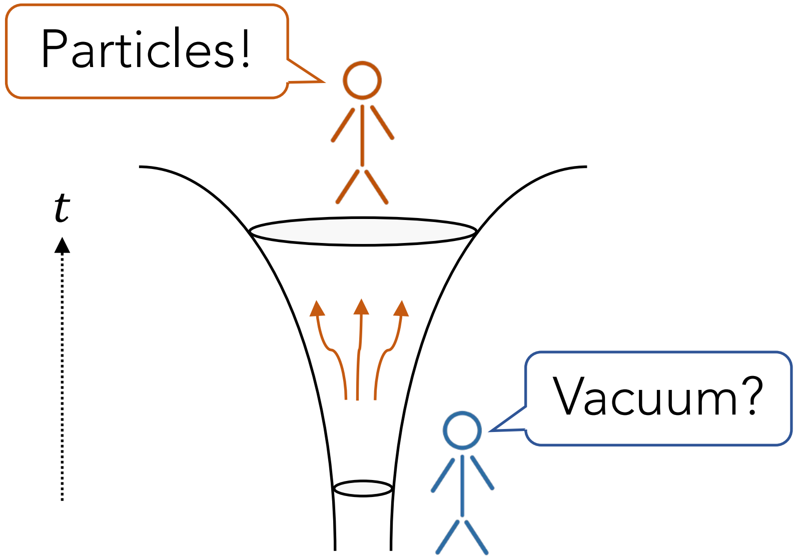
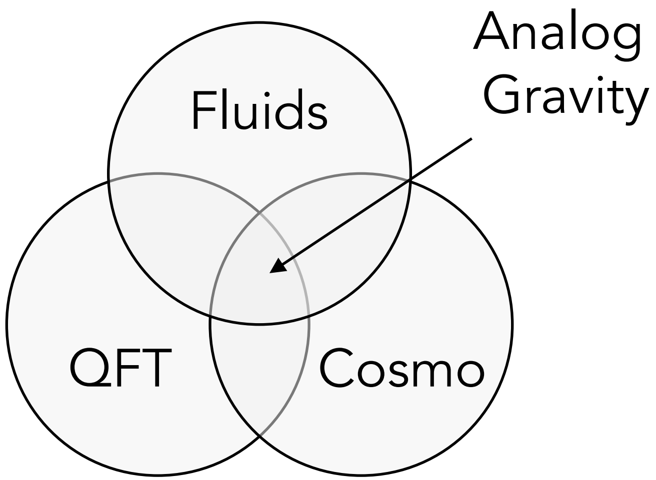
Big picture

- Quantum fields in expanding universe Birrel et al. '82, Mukhanov et al. '07



Cosmological particle production



Cosmological problems ⟷ Fluids

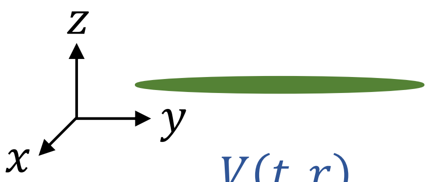
→ Study analog cosmological particle production Jain et al. '07, Eckel et al. '18

Outline

Theory	Experiment
• Mapping BEC $\leftrightarrow$ Universe	• Phonon propagation
• Cosmological particle production	• Density fluctuations

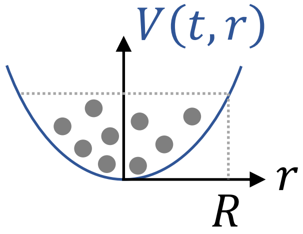
Two-dimensional BEC

- Experimental setup
  - Pancake trap



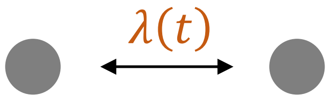
→ 2D geometry

  - Isotropic trap



→ $V(t,r) = \frac{m}{2}\omega^2(t) f(r)$

  - Interactions



→ $\lambda(t) \propto a_s(t)$

- Effective action of Bose-Einstein condensate (BEC) Gross '61, Pitaevskii '61

$$\Gamma[\Phi] = \int dt\, d^2r \left\{ \hbar\Phi^* \left[i\frac{\partial}{\partial t} - V(t,r) \right] \Phi - \frac{\hbar^2}{2m} (\vec{\nabla}\Phi^*)(\vec{\nabla}\Phi) - \frac{\lambda(t)}{2} (\Phi^*\Phi)^2 \right\}$$

Dynamics of Phonons

- Split into background and fluctuations

$$\Phi(t,\vec{r}) \to \phi_0(t,\vec{r}) + \frac{1}{2}[\phi_1(t,\vec{r}) + i\phi_2(t,\vec{r})]$$

- Effective action of phonons ($\phi = \phi_2/\sqrt{2m}$)

$$\Gamma[\phi] = -\frac{\hbar^2}{2}\int dt\, d^2r\, \sqrt{g}\, g^{\mu\nu} \frac{\partial}{\partial x^\mu} \phi \frac{\partial}{\partial x^\nu} \phi$$

in acoustic approx. $\omega(k) = c\,k$ and for $\vec{v}_0 = \text{const.}$

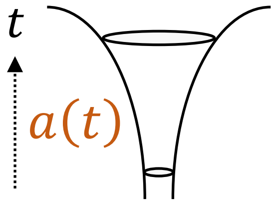
→ Phonons ↔ Relativistic, free, massless, scalar field in curved spacetime

Acoustic metric vs. FLRW metrics

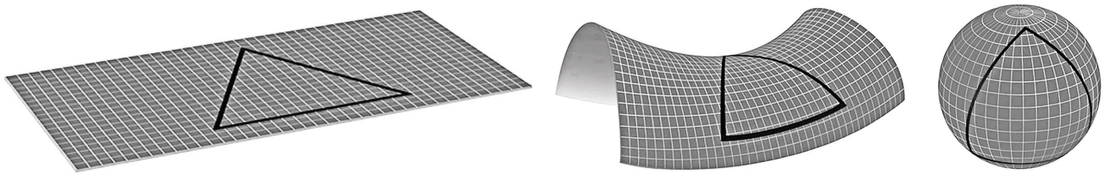
- Spacetime geometry determined by
  - acoustic metric $g_{\mu\nu} = \frac{1}{c^2}\begin{pmatrix} -(c^2 - v_0^2) & -\vec{v}_0 \\ -\vec{v}_0 & 1 \end{pmatrix}$
  - speed of sound $c^2(t,\vec{r}) = \frac{\lambda(t)\, n_0(t,\vec{r})}{m}$ ⟵ time- and space-dependent!
- Friedmann-Lemaître-Robertson-Walker (FLRW) metrics Weinberg '08

$$ds^2 = g_{\mu\nu}dx^\mu dx^\nu = -c^2dt^2 + a^2(t)\left(\frac{du^2}{1-\kappa u^2} + u^2 d\Omega^2\right)$$

Scale factor



Spatial curvature



$\kappa = 0$ $\kappa < 0$ $\kappa > 0$

Engineering expansion

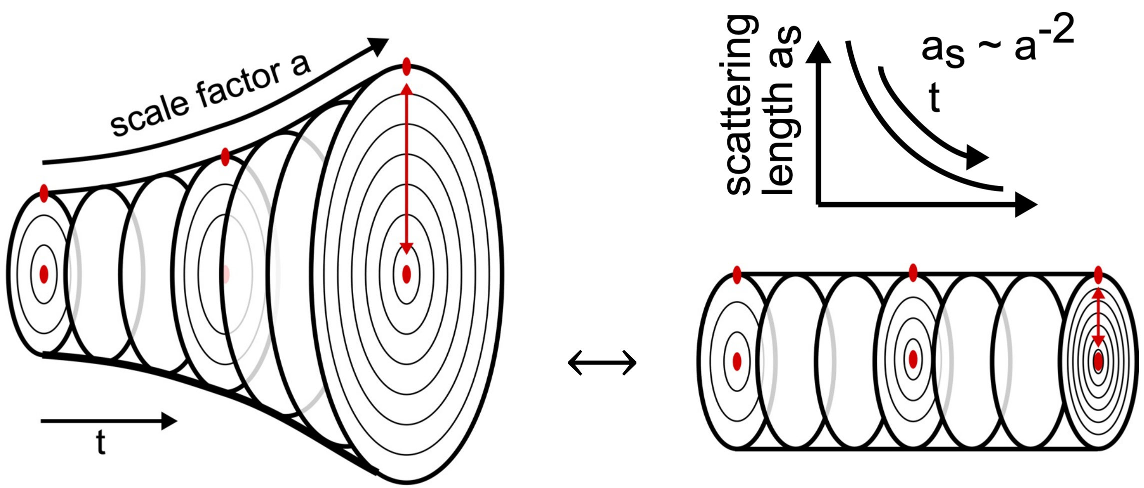
- For scale factor $a(t)$ Jain et al. '07
  - Just define

$$a^2(t) = \frac{m}{\bar{n}_0}\frac{1}{\lambda(t)} \propto \frac{1}{\bar{c}^2(t)}$$

  - and note the equivalence

    expanding space ↔

    decreasing causal speed



→ Coupling $\lambda(t)$ ↔ Scale factor $a(t)$

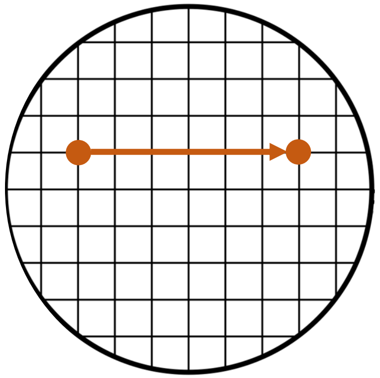
Engineering curvature - I

- For flat universe $\kappa = 0$

Density profile

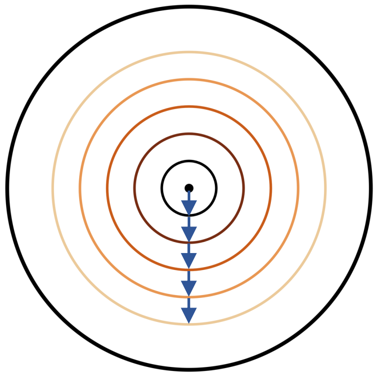
$$n_0(r) = \bar{n}_0$$

Geometry



Geodesics:
straight

Phonons



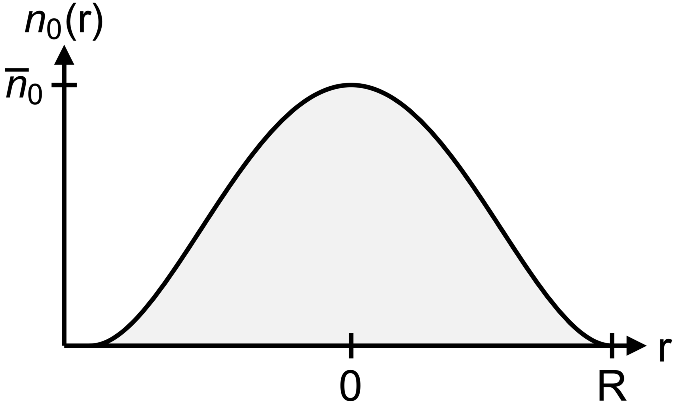
Radial distance:
constant

Engineering curvature - II

Mapping BEC ↔ Universe | Phonon propagation | Cosmological particle production | Density fluctuations

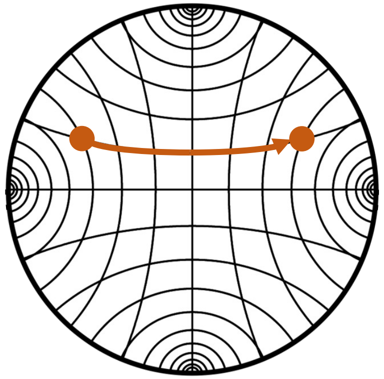
- For hyperbolic universe $\kappa < 0$

Density profile



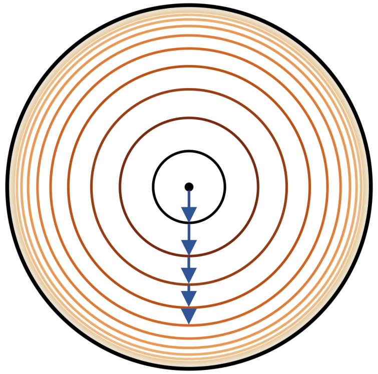
$$n_0(r) = \bar{n}_0 \left(1 - \frac{r^2}{R^2}\right)^2$$

Geometry



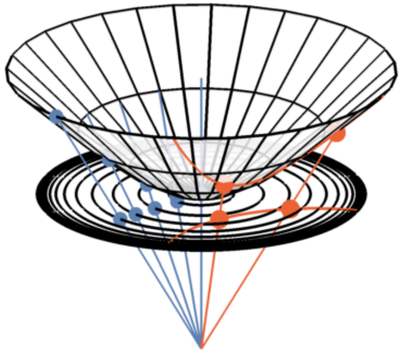
Geodesics:
arcs + dia.

Phonons



Radial distance:
decreases

Projection

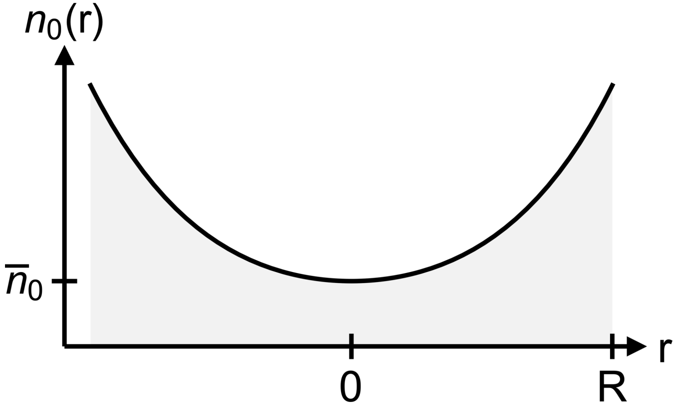


Poincaré disc

Engineering curvature - III

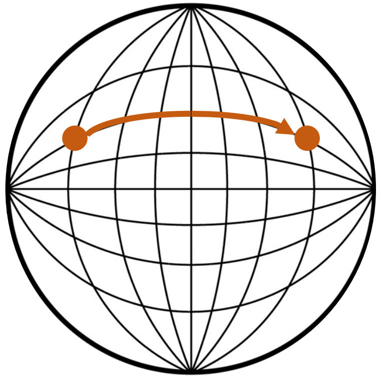
- For spherical universe $\kappa > 0$

Density profile



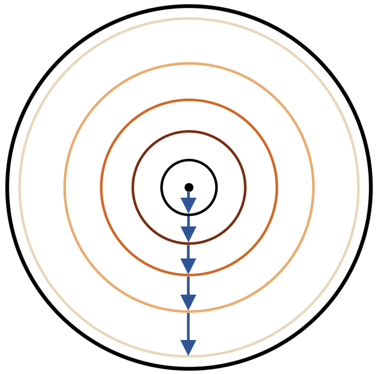
$$n_0(r) = \bar{n}_0 \left(1 + \frac{r^2}{R^2}\right)^2$$

Geometry



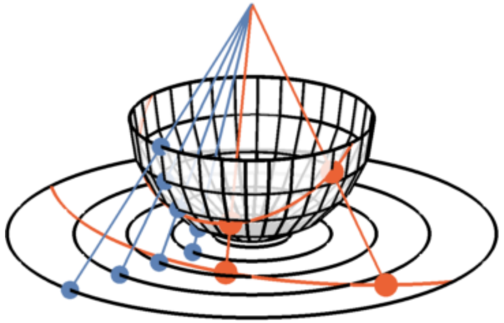
Geodesics:
arcs + dia.

Phonons



Radial distance:
increases

Projection

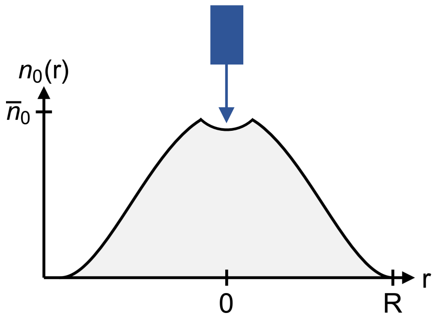


Stereographic disc

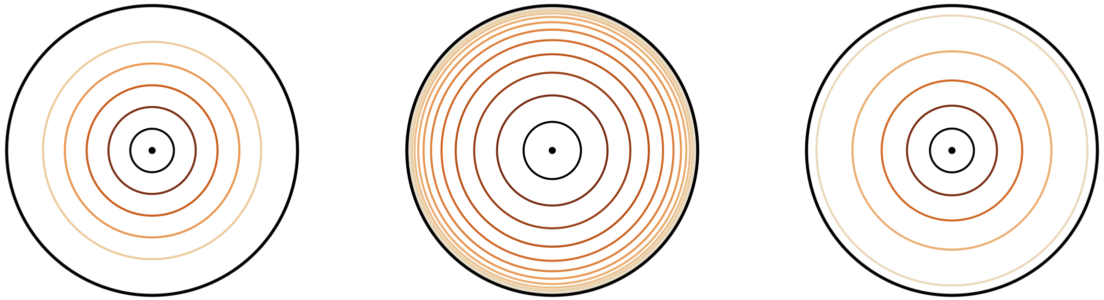
→ Density profile $n_0(r)$ ↔ Spatial curvature κ

Imprinting phononic waves

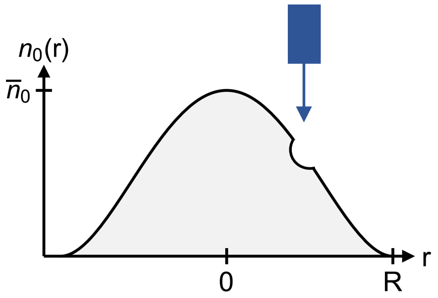
- Perturb condensate with a laser
  - In the center



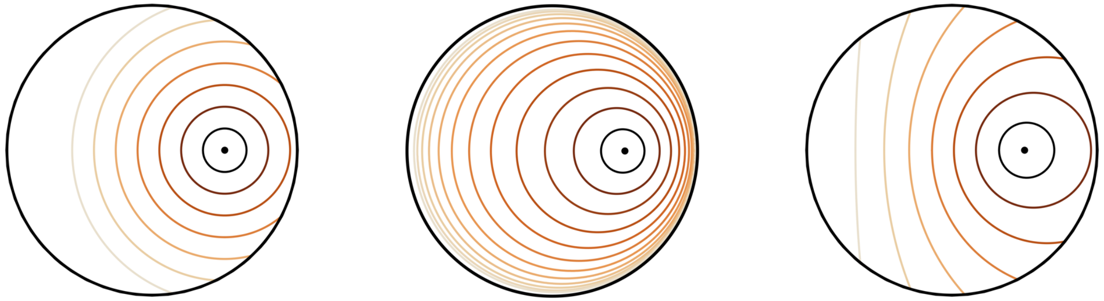
$\kappa = 0$ $\kappa < 0$ $\kappa > 0$



  - Somewhere else

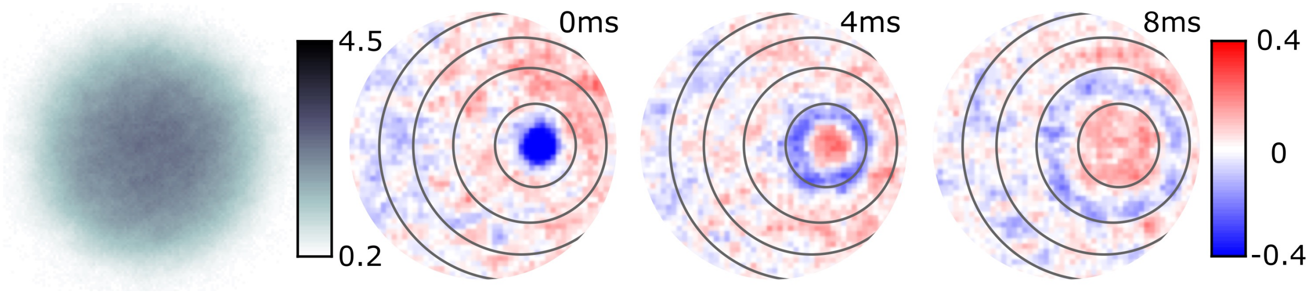


$\kappa = 0$ $\kappa < 0$ $\kappa > 0$

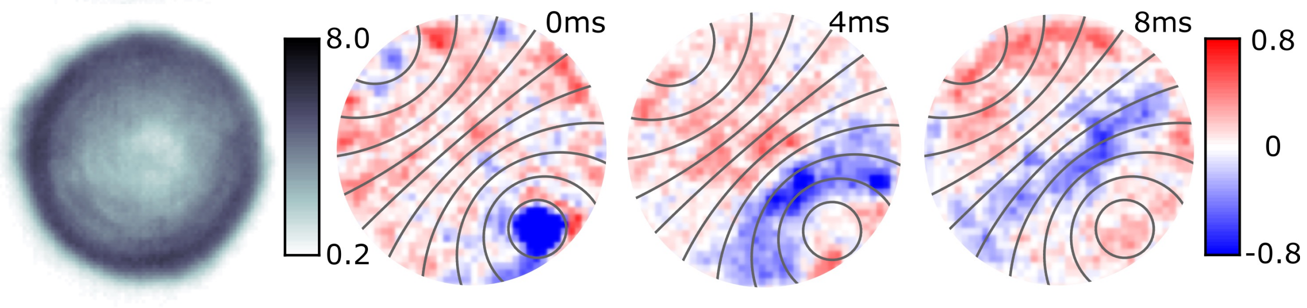


Configurability of spatial curvature

- Hyperbolic geometry ($\kappa < 0$)

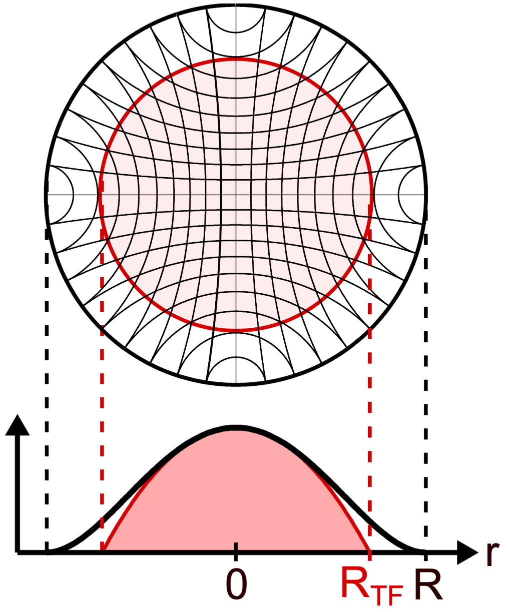


- Spherical geometry ($\kappa > 0$)



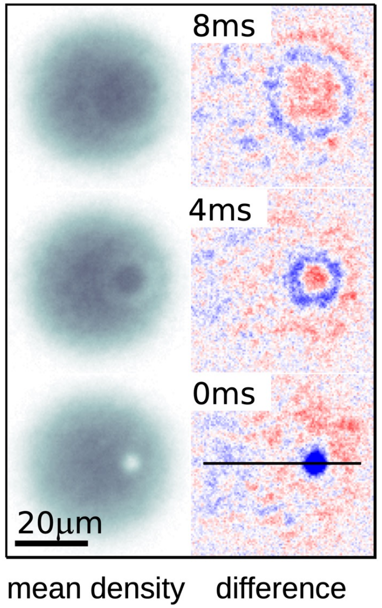
Phonon trajectories

- Hyperbolic geometry ($\kappa < 0$)

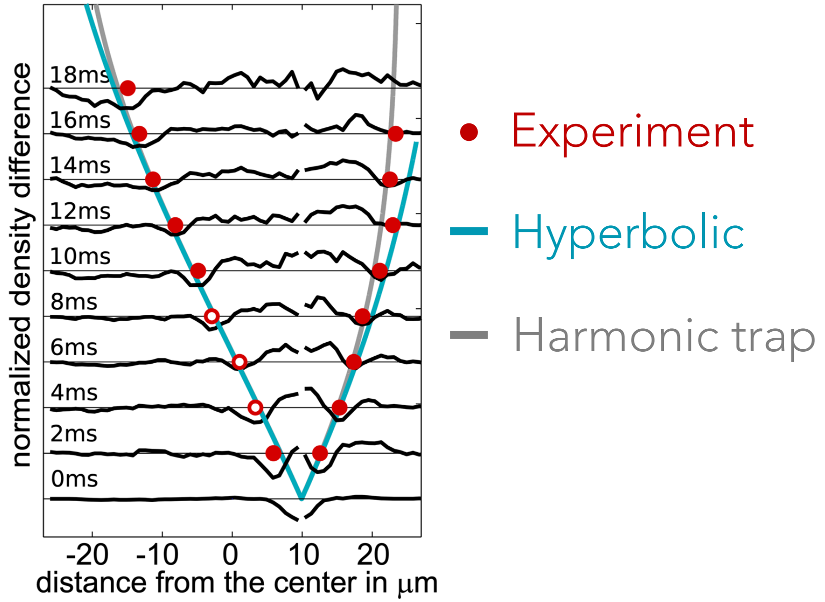


Harmonic trap

$R = \sqrt{2}\, R_{TF}$



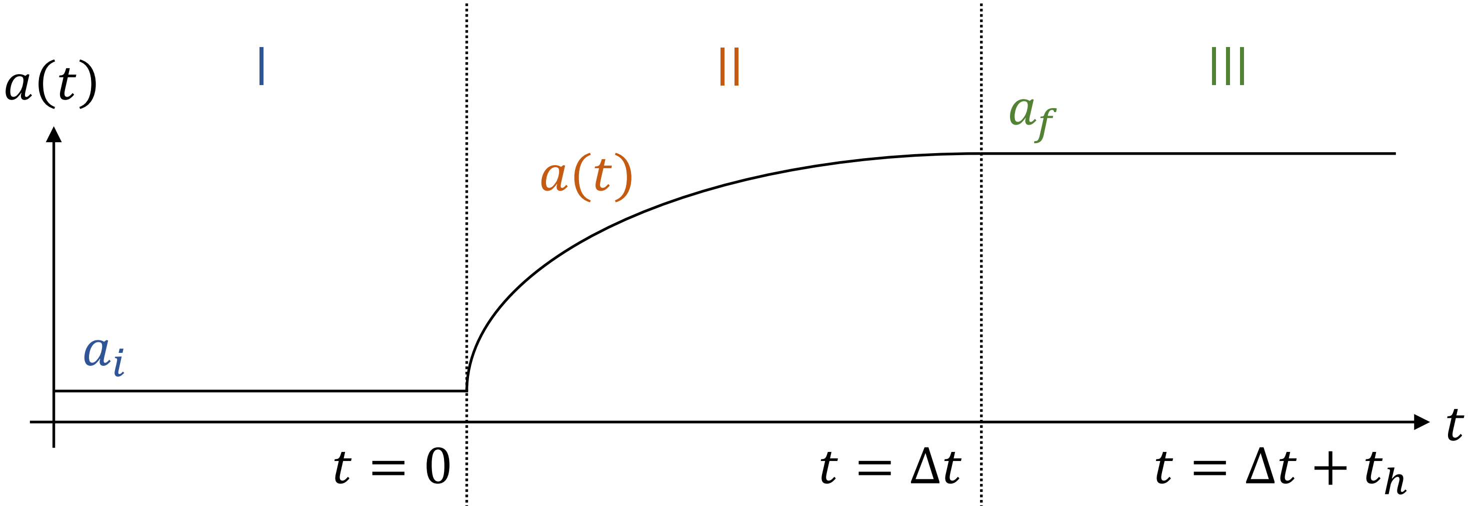
Phonon propagation



Quantitative comparison

Expansion and quanta

- Controlled expansion Jain et al. '07
  - Regions



  - Particles/Vacua: $\hat{a}_{km}|\Omega\rangle = 0$ | ambiguous | $\hat{b}_{km}|\Psi\rangle = 0$

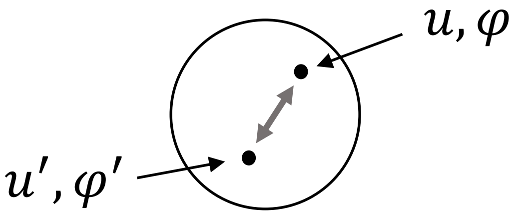
- Initial vacuum $|\Omega\rangle$ is not empty anymore $\hat{b}_{km}|\Omega\rangle \neq 0$ → Particle production

Correlation functions and spectra

- Rescaled density contrast

$$\delta_c(t,u,\varphi) = \sqrt{\frac{n_0(u)}{\bar{n}_0^3}}\,[n(t,u,\varphi) - n_0(u)]$$

- Correlations between two points after expansion $t \geq \Delta t$



$$\langle \delta_c(t,u,\varphi)\delta_c(t,u',\varphi')\rangle = \text{const.} \times \langle \dot{\phi}\dot{\phi}\rangle(t,L)$$

$$= \text{const.} \times \int_k \mathcal{F}(k,L)\sqrt{-h(k)}\,S_k(t)\,\tilde{f}_G(k)$$

with spectrum of fluctuations

$$S_k(t) = \frac{1}{2} + N_k + \Delta N_k(t) = \frac{1}{2} + |\beta_k|^2 + |\alpha_k\beta_k|\cos(2\omega_k t + \Theta_k)$$

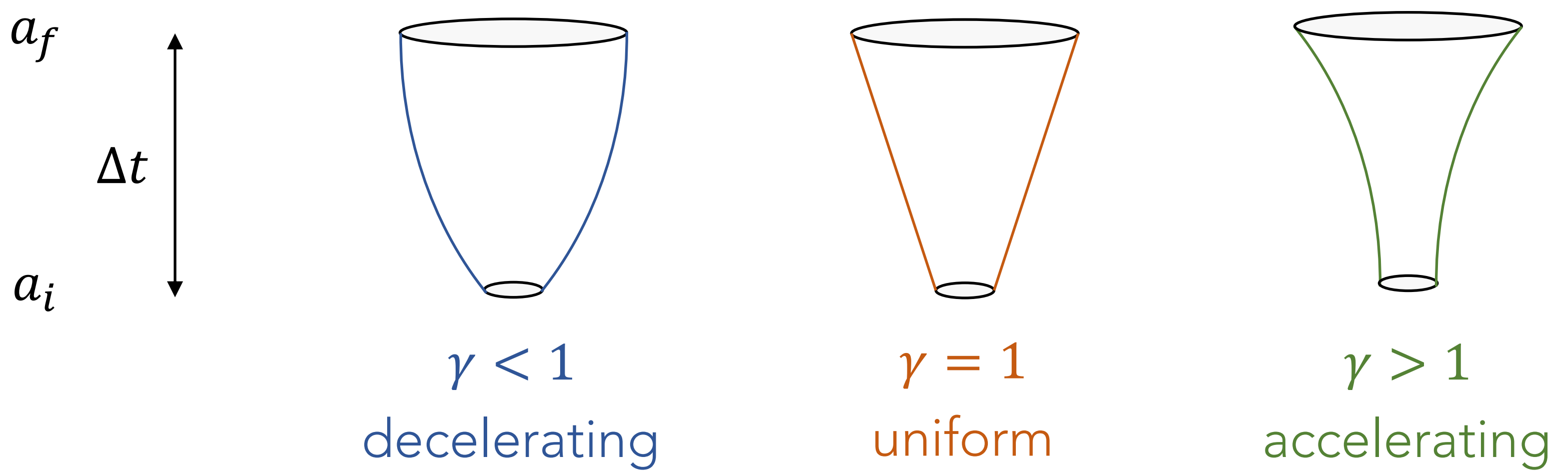
Bogoliubov coefficients: $\hat{b}_{km} = \alpha_k^*\hat{a}_{km} - \beta_k^*(-1)^m\hat{a}_{k,-m}^\dagger$

Expansion scenarios

- Family of power-law expansions

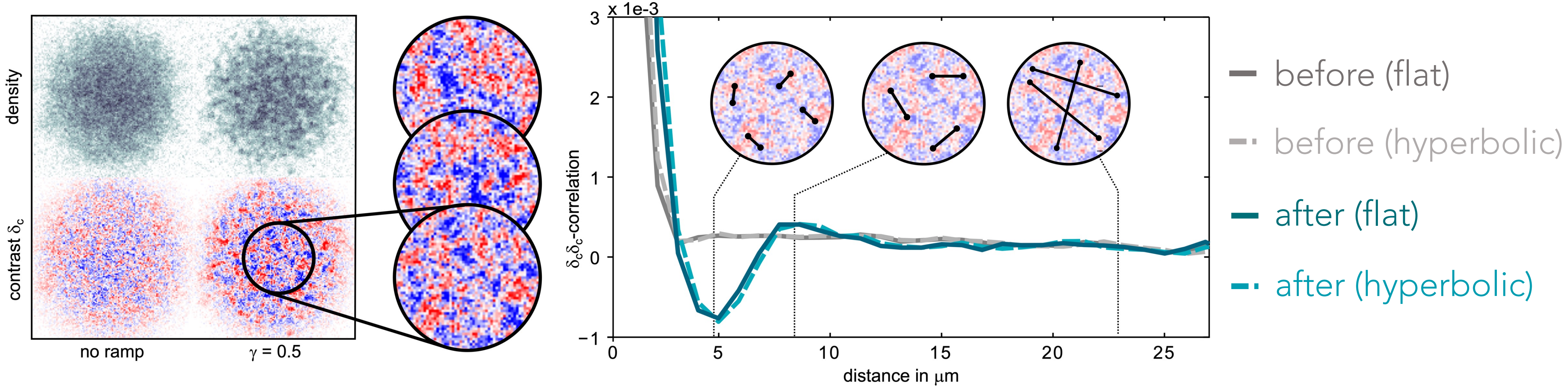
$$a(t) = \text{const.} \times |t - t_0|^{\gamma}$$

- Fix initial/final sizes a_i, a_f and expansion duration Δt → Three expansion classes



Extracting correlation functions

- Examples for $\gamma = 0.5$ (decelerating) and $\kappa < 0$ (harmonic trap) at $t = \Delta t = 1.5$ ms

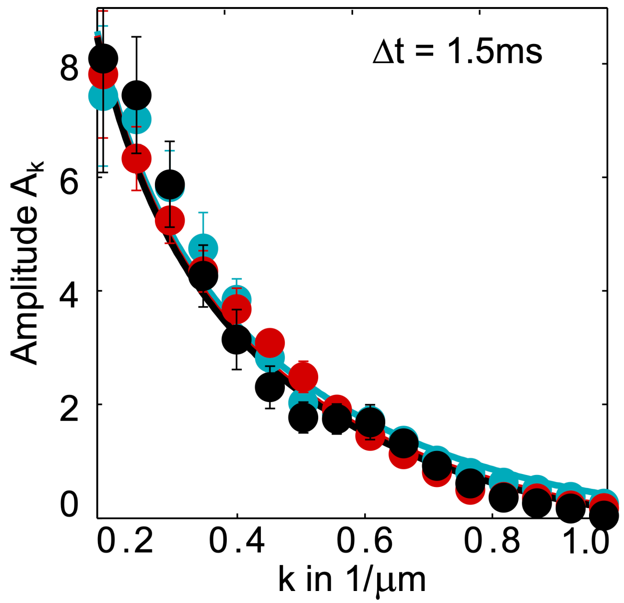


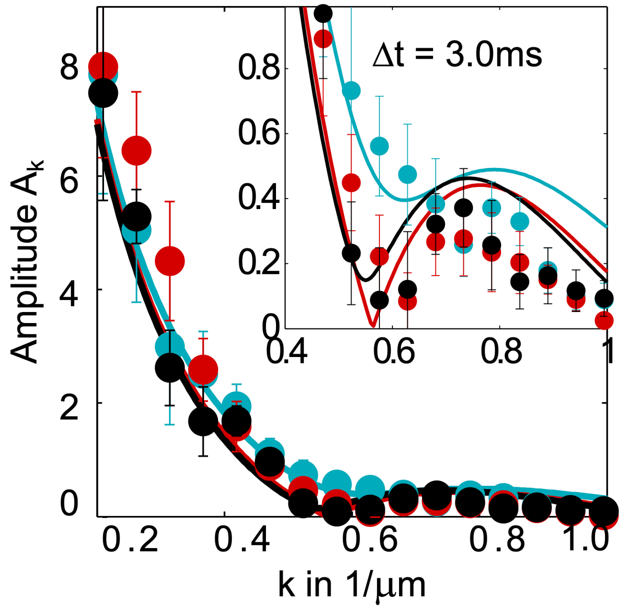
single realization | inner region | averaged over 100 realizations and all distances L

→ Clear signal and flat approximation works quite well in the center

Spectra: Amplitudes

- For $\kappa < 0$ (harmonic trap)





$\gamma = 0.5$

$\gamma = 1$

$\gamma = 1.5$

• Experiment

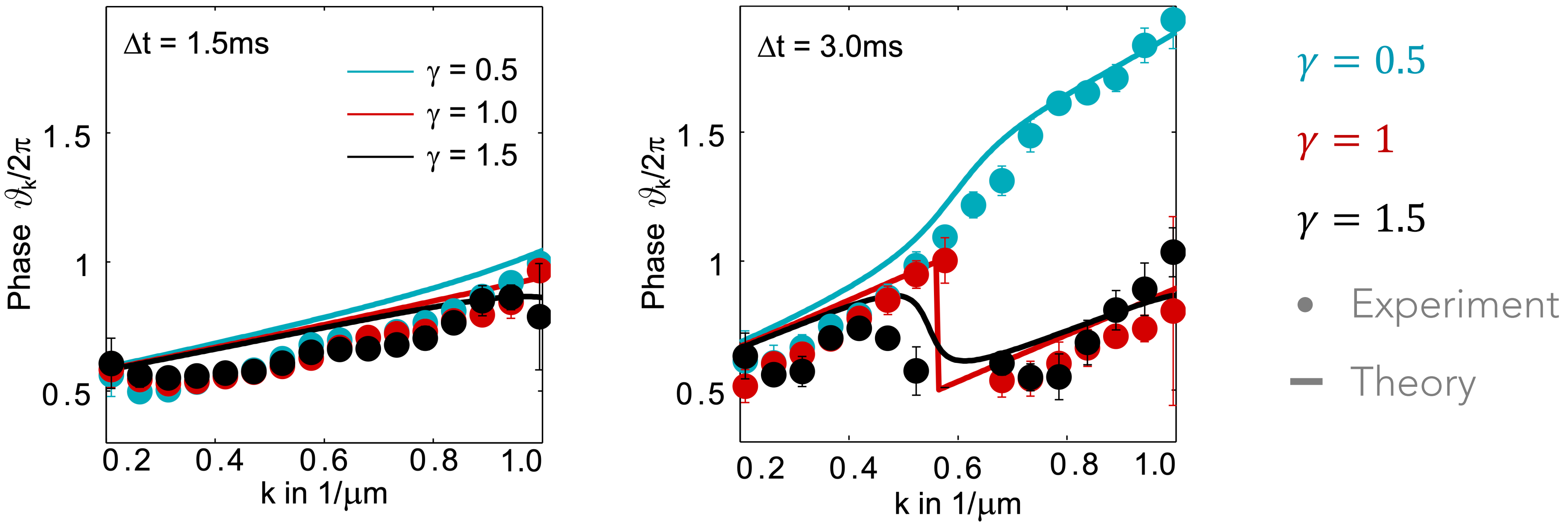
— Theory

→ Dependence on initial state

→ Hard to distinguish between expansion scenarios

Spectra: Phases

- For $\kappa < 0$ (harmonic trap)



→ Phases are robust and have decisive features

→ Analog cosmological particle production confirmed

Summary

Theory

- Mapping BEC $\leftrightarrow$ Universe

Phonons	$\leftrightarrow$	Relativistic scalar field
Density profile $n_0(r)$	$\leftrightarrow$	Spatial curvature κ
Coupling $\lambda(t)$	$\leftrightarrow$	Scale factor $a(t)$

- Cosmological particle production

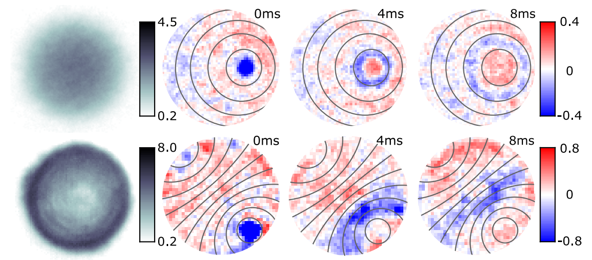
before — expansion $a(t)$ — after

$$\hat{a}_{km}|\Omega\rangle = 0 \xrightarrow[\alpha_k,\beta_k]{} \hat{b}_{km}|\Psi\rangle = 0$$

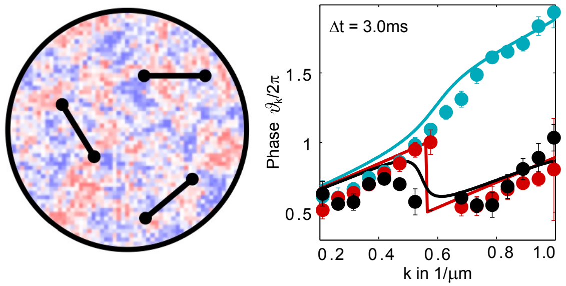
$$S_k(t) = \frac{1}{2} + |\beta_k|^2 + |\alpha_k\beta_k| \cos(2\omega_k t + \Theta_k)$$

Experiment

- Phonon propagation



- Density fluctuations

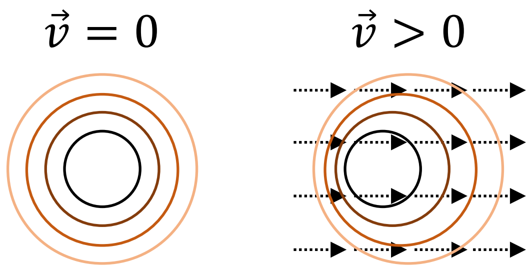


Thank you for your attention :)

Backup

A simple analogy

- Sound waves in a fluid[Novello et al. '02], which is described by



- Continuity equation $\frac{\partial \rho}{\partial t} + \vec{\nabla}(\rho \vec{v}) = 0,$
- Euler equation $\rho\left(\frac{\partial \vec{v}}{\partial t} + (\vec{v}\cdot\vec{\nabla})\vec{v}\right) = -\nabla p$

- Sound waves: Fluctuations on background $\rho \to \rho_0 + \rho_1, p \to p_0 + p_1, \phi \to \phi_0 + \phi_1$
- E.o.m. of sound waves $\frac{1}{\sqrt{g}}\frac{\partial}{\partial x^\mu}\left(\sqrt{g}\, g^{\mu\nu}\frac{\partial}{\partial x^\nu}\ \phi_1\right) = 0$

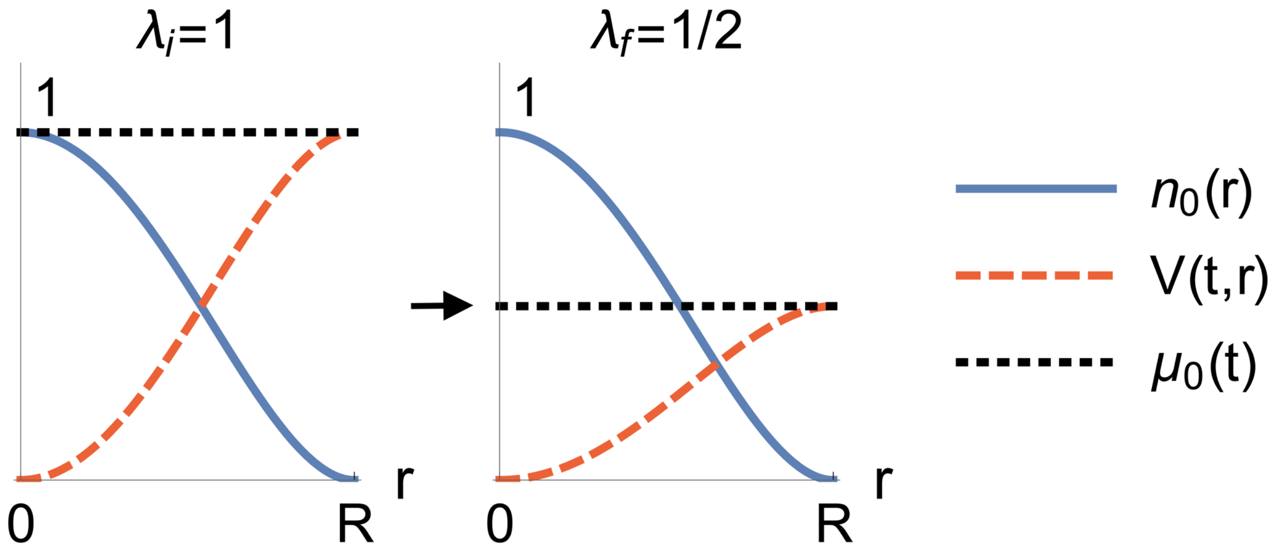
acoustic metric $g_{\mu\nu} \propto \begin{pmatrix} -(c^2 - v_0^2) & -\vec{v}_0 \\ -\vec{v}_0 & 1 \end{pmatrix}$, speed of sound $c^2 = \frac{\partial p}{\partial \rho}$

→ Sound waves in a fluid ⟷ Scalar field in curved spacetime

Keeping background density static

- Changing the coupling $\lambda(t)$ is accompanied by

$$\omega^2(t) = \frac{2\bar{n}_0}{mR^2}\,\lambda(t), \qquad \mu_0(t) = \bar{n}_0\,\lambda(t)$$



More details on the mapping

- Acoustic line element for appropriate density profiles $n_0(r)$

$$ds^2 = -dt^2 + \frac{m}{\bar{n}_0}\frac{1}{\lambda(t)}\left(1 \mp \frac{r^2}{R^2}\right)^{-2}(dr^2 + r^2 d\varphi^2)$$

- Define

$$a^2(t) = \frac{m}{\bar{n}_0}\frac{1}{\lambda(t)}, \qquad u(r) = \frac{r}{1 \mp \frac{r^2}{R^2}} \quad \rightarrow \quad \frac{dr^2}{\left(1 \mp \frac{r^2}{R^2}\right)^2} = \frac{du^2}{1 \pm 4\frac{u^2}{R^2}}$$

- Arrive at FLRW metric

$$ds^2 = -dt^2 + a^2(t)\left(\frac{du^2}{1 - \kappa u^2} + u^2 d\varphi^2\right)$$

with $\kappa = \mp 4/R^2$

Quantization of phonon field

- E.o.m. of phonon field ϕ with FLRW metric $g_{\mu\nu}$

$$\ddot{\phi} + 2\frac{\dot{a}(t)}{a(t)}\dot{\phi} - \frac{1}{a^2(t)}\Delta\phi = 0$$

- Promote phonon field ϕ to a quantum operator $\hat{\phi}$

$$\hat{\phi}(t,u,\varphi) = \int_{k,m} \left[\hat{a}_{km}\,\mathcal{H}_{km}(u,\varphi)\,v_k(t) + \hat{a}^\dagger_{km}\,\mathcal{H}^*_{km}(u,\varphi)\,v^*_k(t)\right]$$

Annihilation op. ($\hat{a}_{km}$)

$$\Delta\mathcal{H}_{km}(u,\varphi) = h(k)\mathcal{H}_{km}(u,\varphi)$$

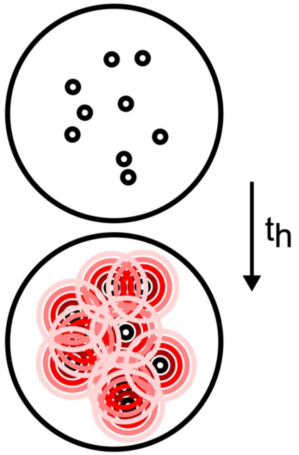
Mode functions ($v_k(t)$)

- Obtain mode equation

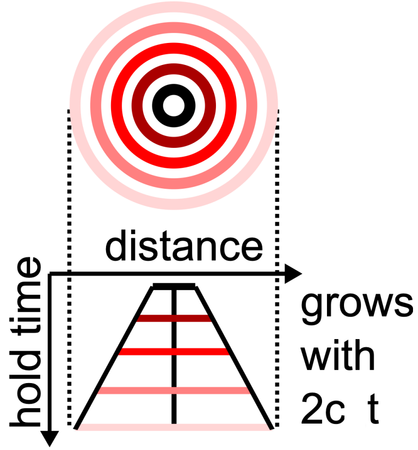
$$\ddot{v}_k(t) + 2\frac{\dot{a}(t)}{a(t)}\dot{v}_k(t) - \frac{h(k)}{a^2(t)}v_k(t) = 0$$

Propagation of correlations

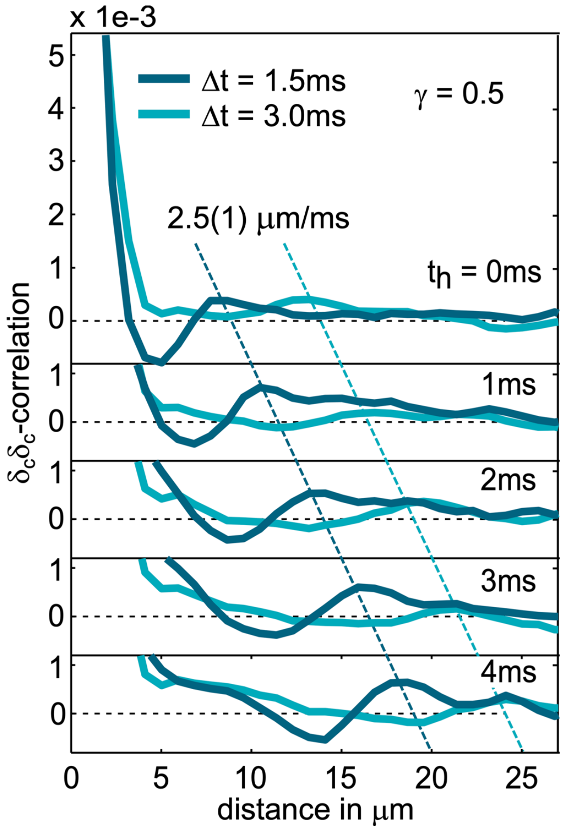
- Again $\gamma = 0.5$ (decelerating) and $\kappa < 0$ (harmonic trap), but now $t = \Delta t + t_h$



excitations
interfer



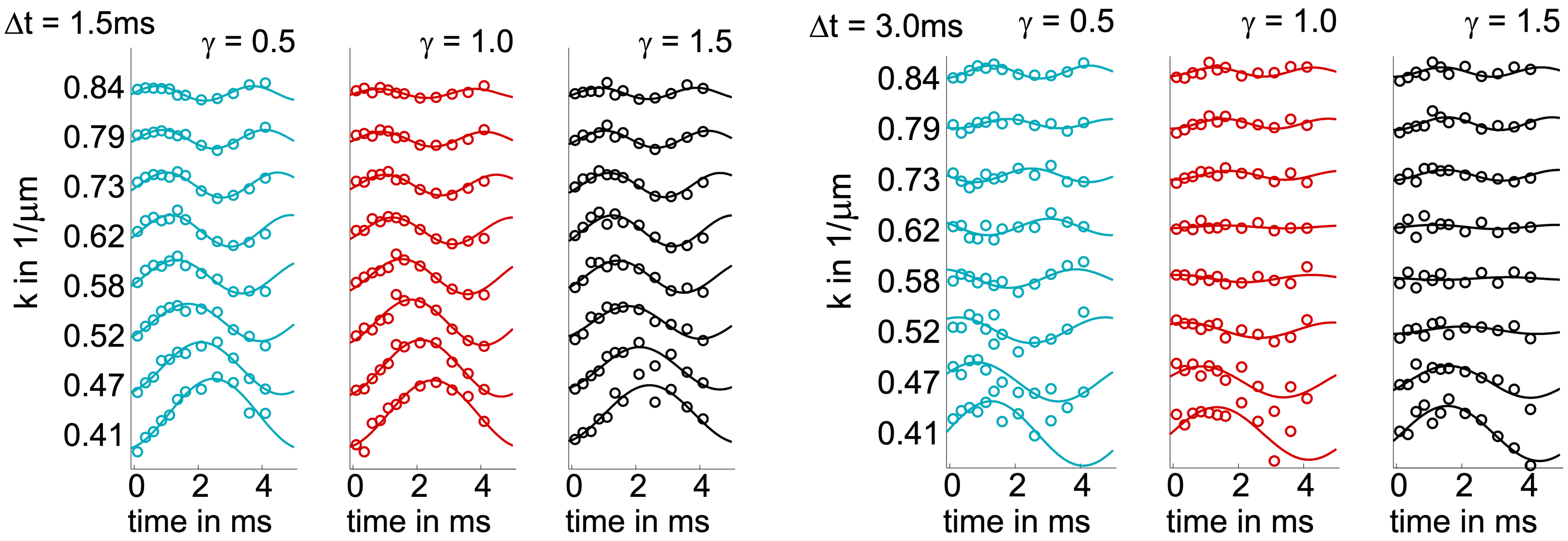
sound cone
grows



→ Correlations propagate with $2c$

Extracting spectra

- Time evolution of k-modes for all three classes and $\kappa < 0$ (harmonic trap)

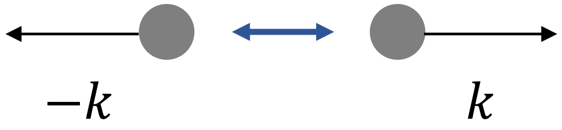


- Fit to $S_k(t) = \frac{1}{2} + |\beta_k|^2 + |\alpha_k \beta_k| \cos(2\omega_k t + \Theta_k)$

→ Extract amplitude $A_k = |\alpha_k \beta_k|$ and phase Θ_k of k-modes

Outlook

- Simulate other cosmologies
  - Vary spatial curvature κ in time
  - Expansions with accelerating ($\gamma > 1$) and decelerating ($\gamma < 1$) epochs
  - Horizons
- Certify quantum nature of produced particles
  - Particles are produced in pairs with opposite momenta



→ Detect entanglement

Experimental values

• Atom species	^{39}K
• Atom number	$\approx 23\,000$
• Trapping frequency in z-direction	$\omega_z = 2\pi \times 1.6$ kHz
• Trapping frequency in r-direction	$\omega_r = 2\pi \times (7 - 23)$ Hz
• Thomas-Fermi radius	$R_{TF} = (25, 30)\ \mu$m
• Scattering length	$a_s = (50 - 400)\ a_B$
• Imaging resolution	$1\ \mu$m